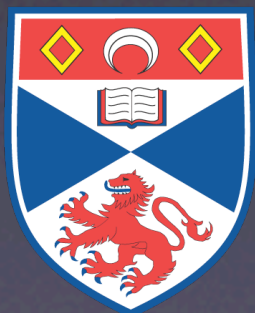


Model SED fitting

(or ‘our bad data extraction algorithm’)

Thomas Robitaille

Barb Whitney, Remy Indebetouw, Kenneth Wood



University
of
St Andrews



Why? (I)

- Model SEDs or spectra are necessary to find synthetic colors (e.g. of YSOs)
- Analyzing how colors depend on various parameters, and comparing real colors to predicted trends is indirect
- Why not compare the models directly to all the data available for each source?

Why? (II)

- Optimal use of available data
- Find not only one set but many sets of parameters which fit data
- Avoid recomputing radiation transfer models for each source

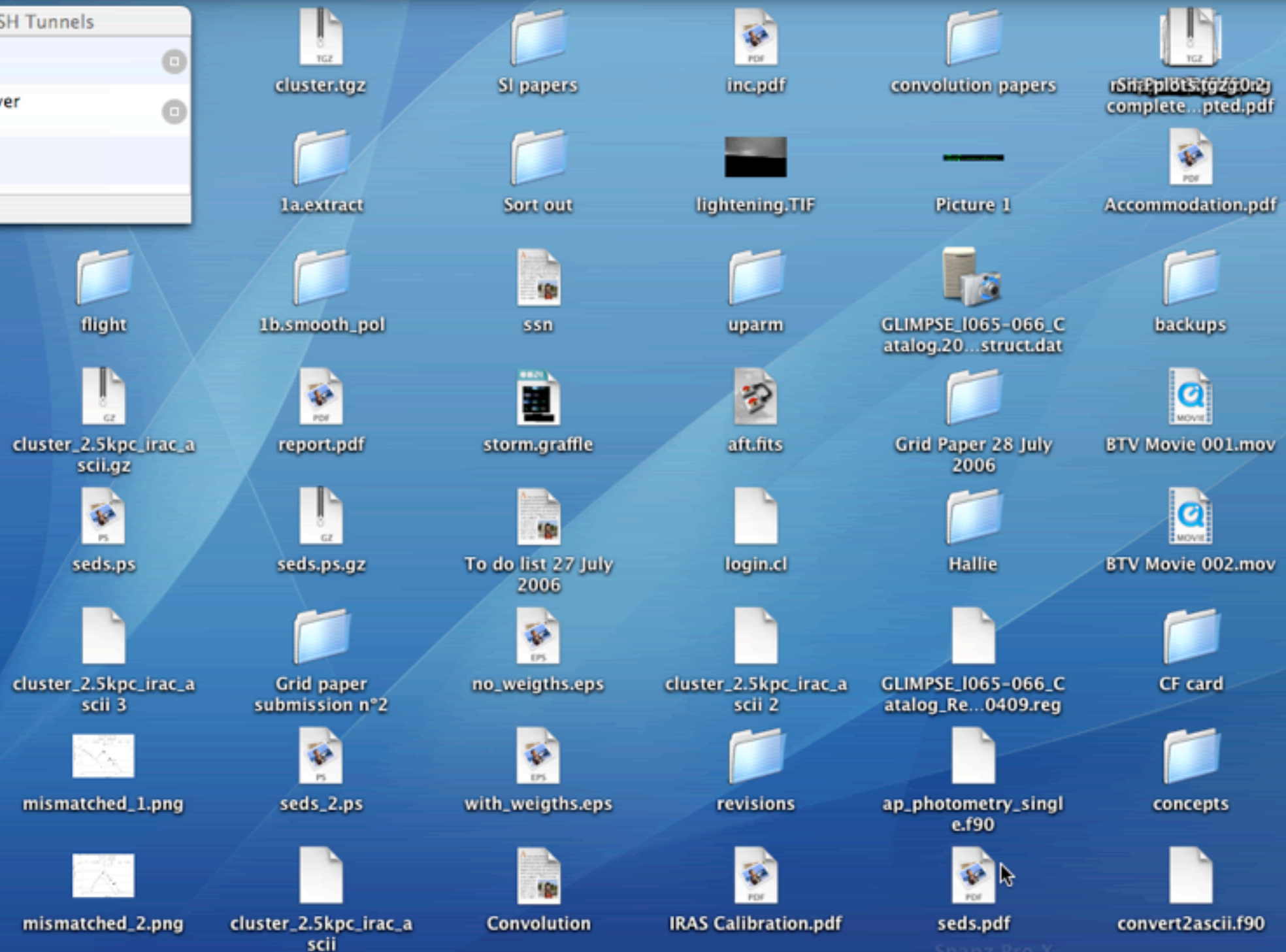
How?

- Use pre-computed model SEDs for YSOs and stars.
 - ➔ 200,000 YSO model SEDs (Robitaille et al 2006)
 - ➔ 7800 PHOENIX stellar photospheres (Brott & Hauschild 2005)
- For each source:
 - ➔ Fit all models to minimize χ^2
 - ➔ Order in decreasing χ^2

SSH Tunnels

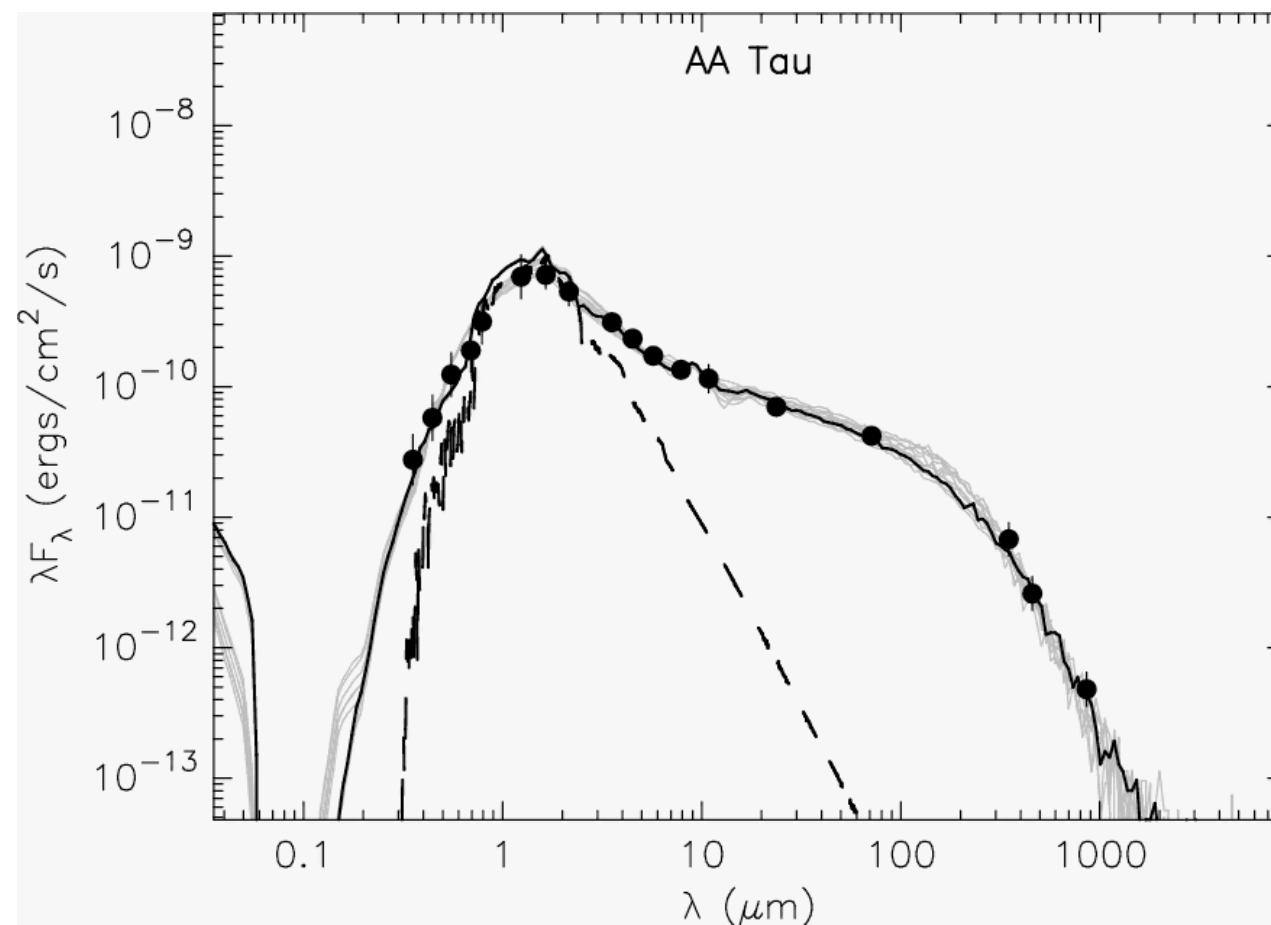
- Uni Email
Connected
- SMTP Server
Connected

Configuration



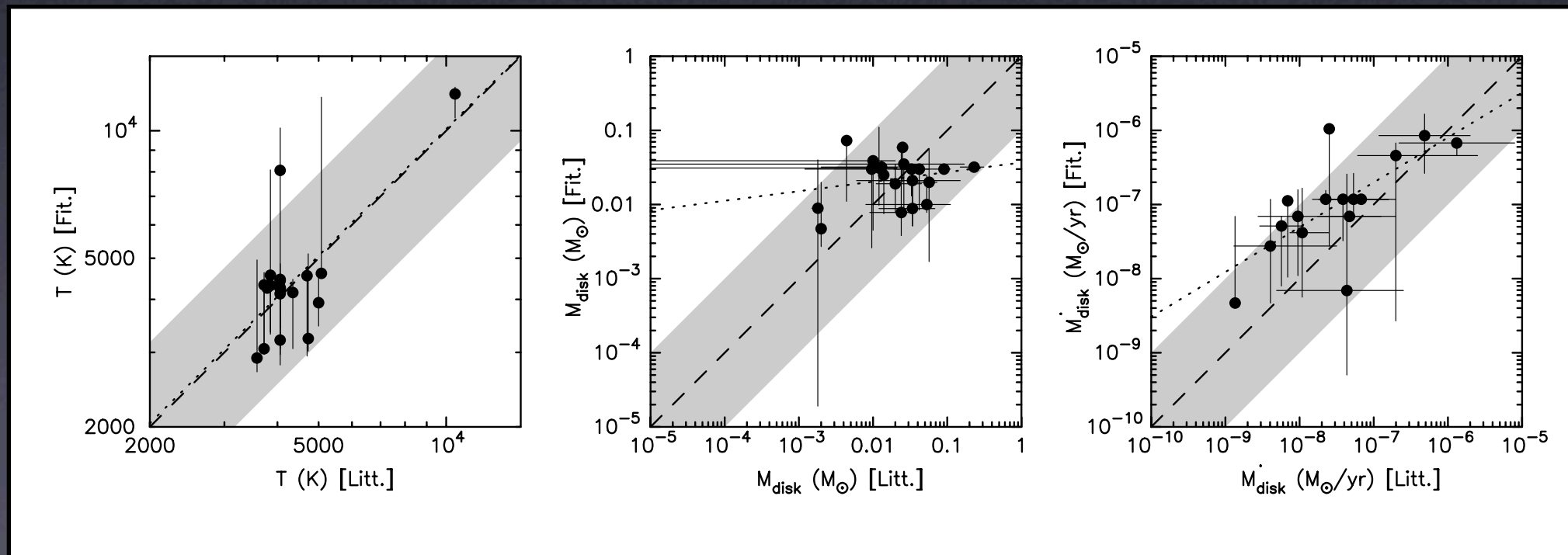
Does it work? (YSOs I)

- To answer this question, we have carried out an analysis of 30 known Taurus sources.

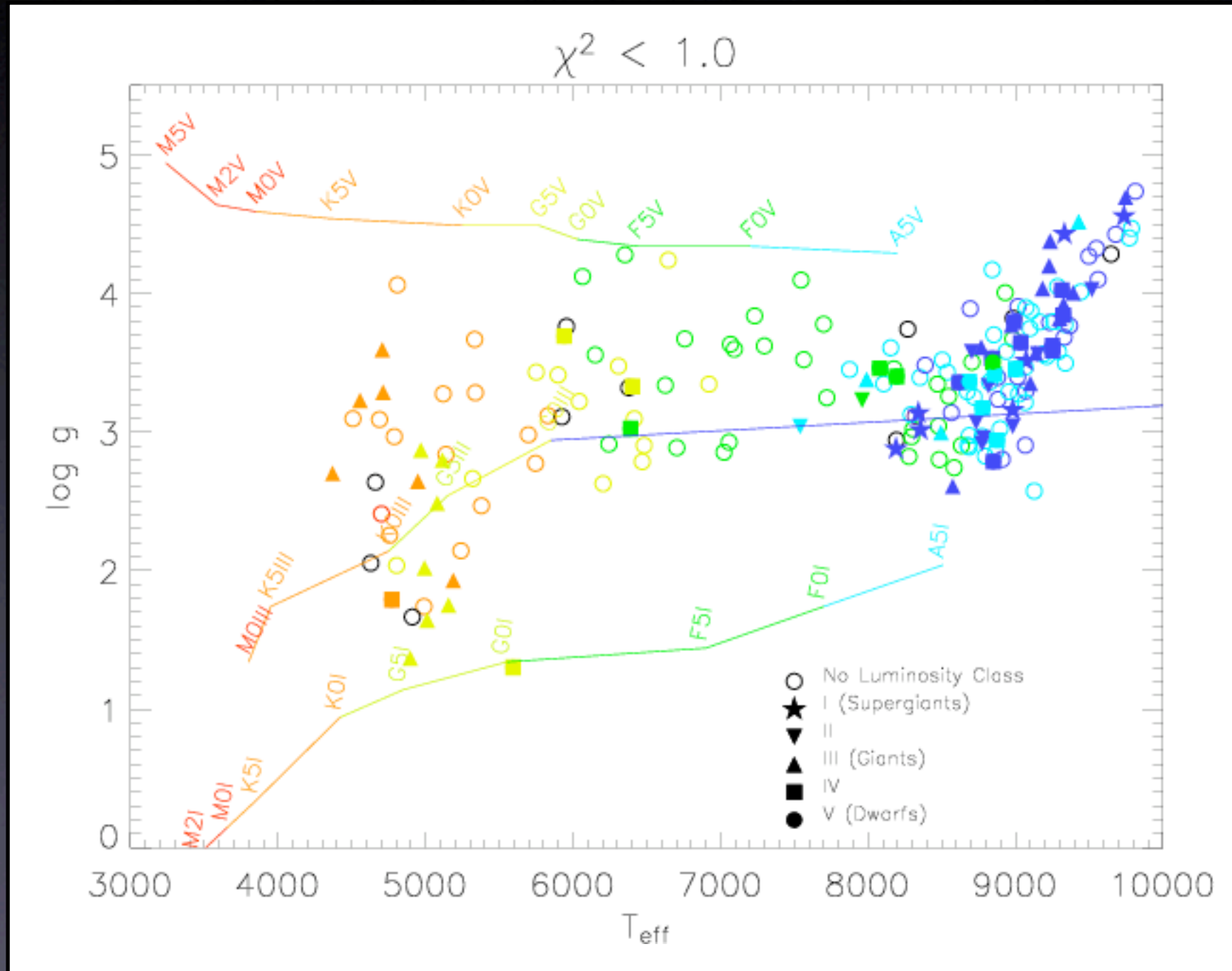


Does it work? (YSOs II)

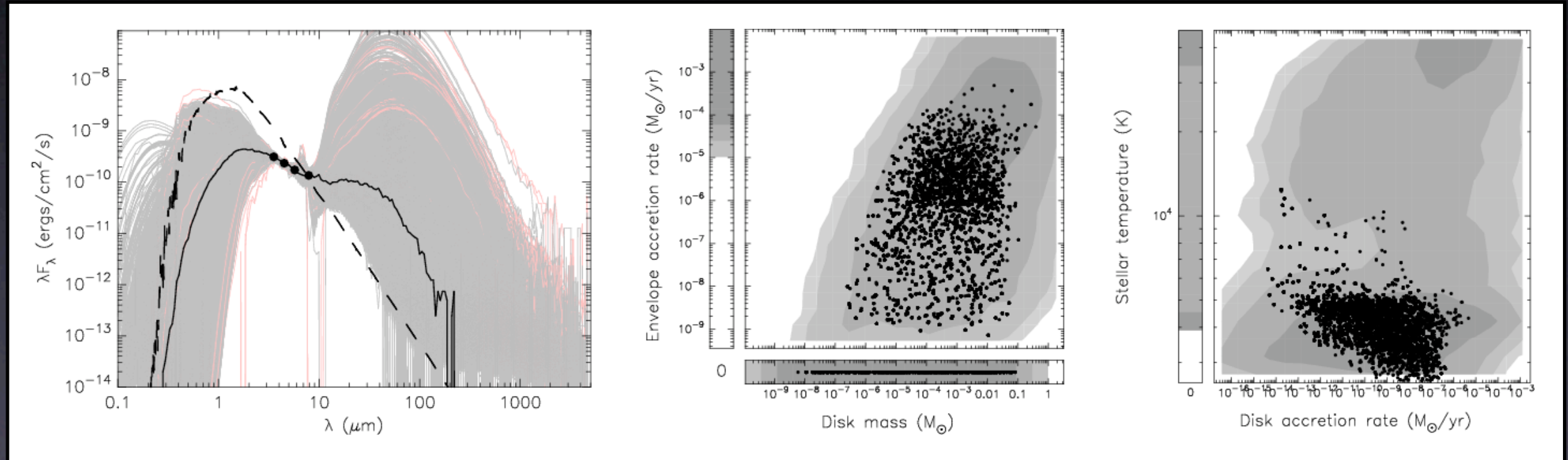
- We compared the parameter values of the good SED fits with independently determined physical parameters:



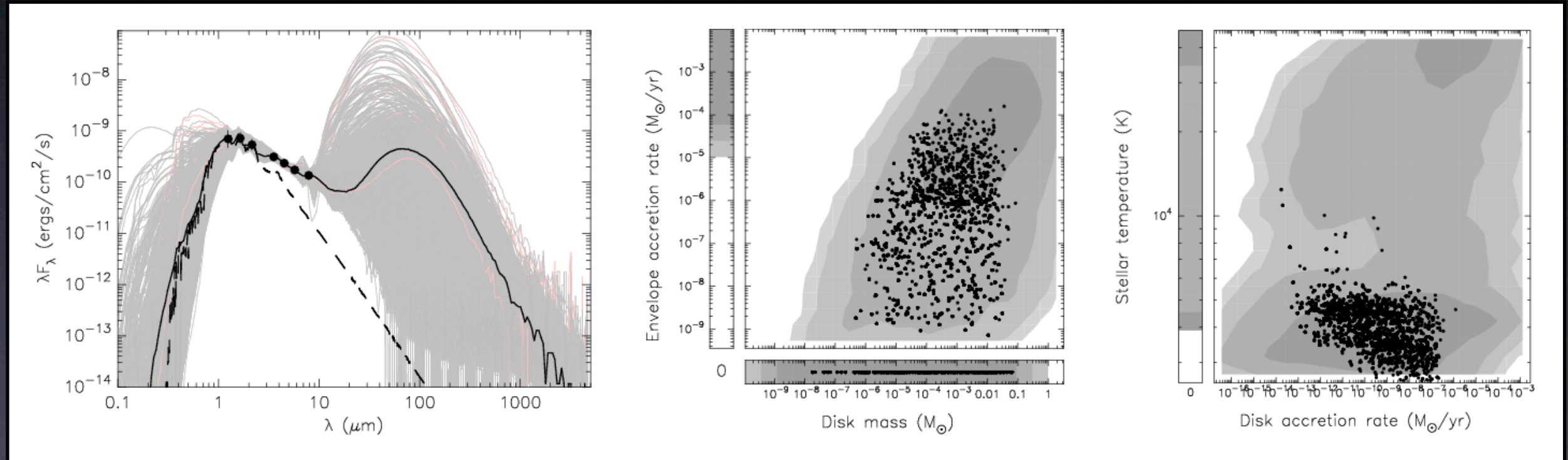
Does it work? (Stars)



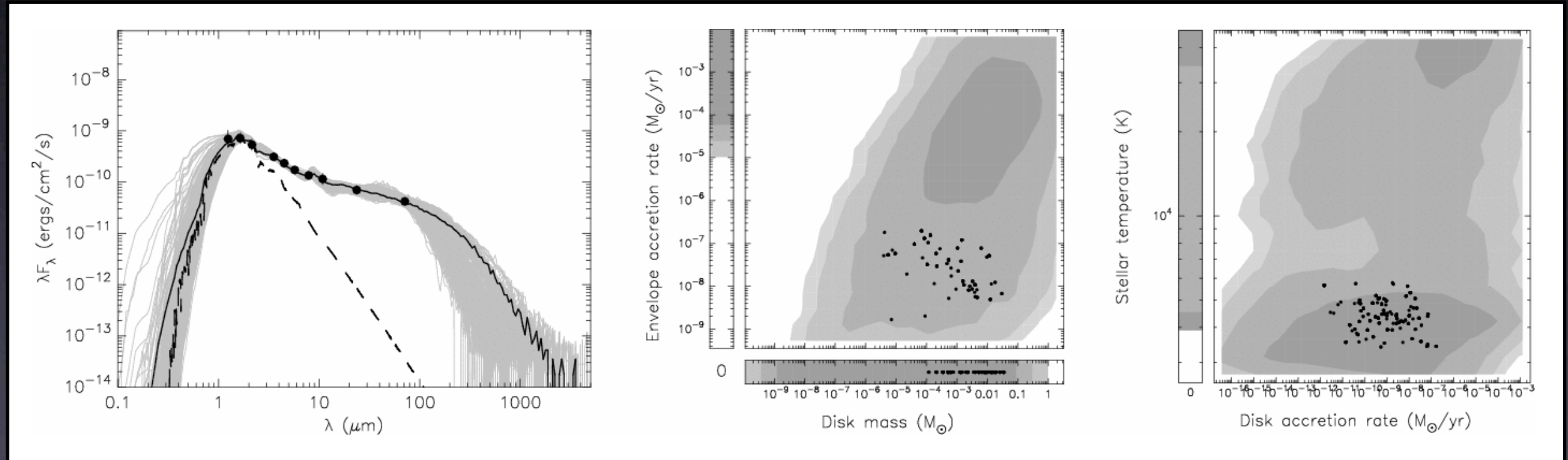
What data do we need?



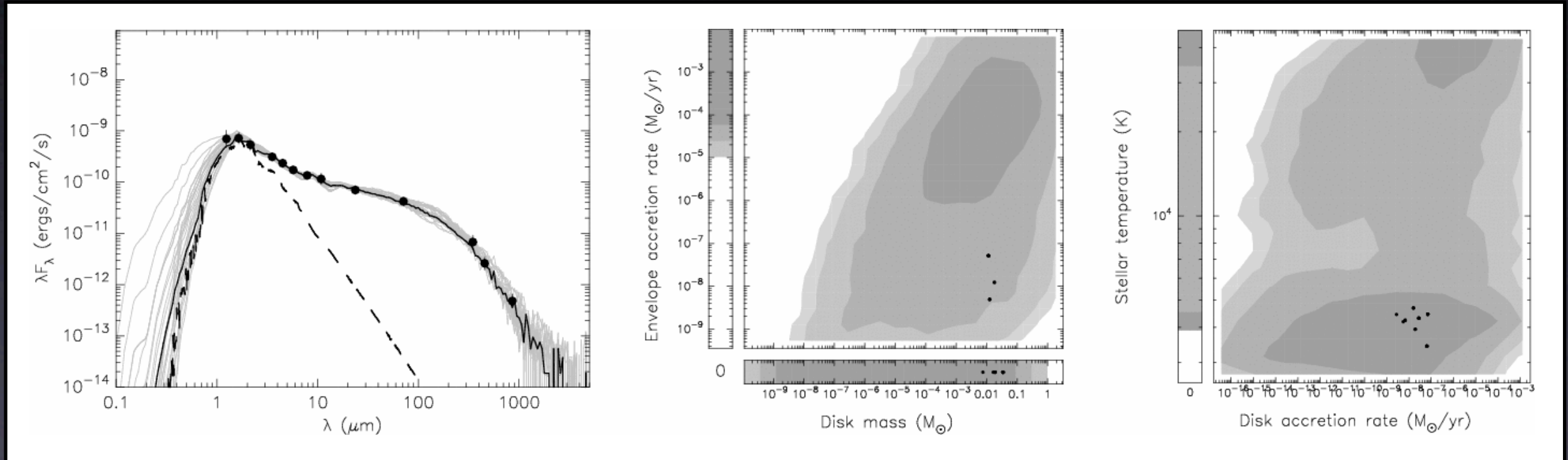
What data do we need?



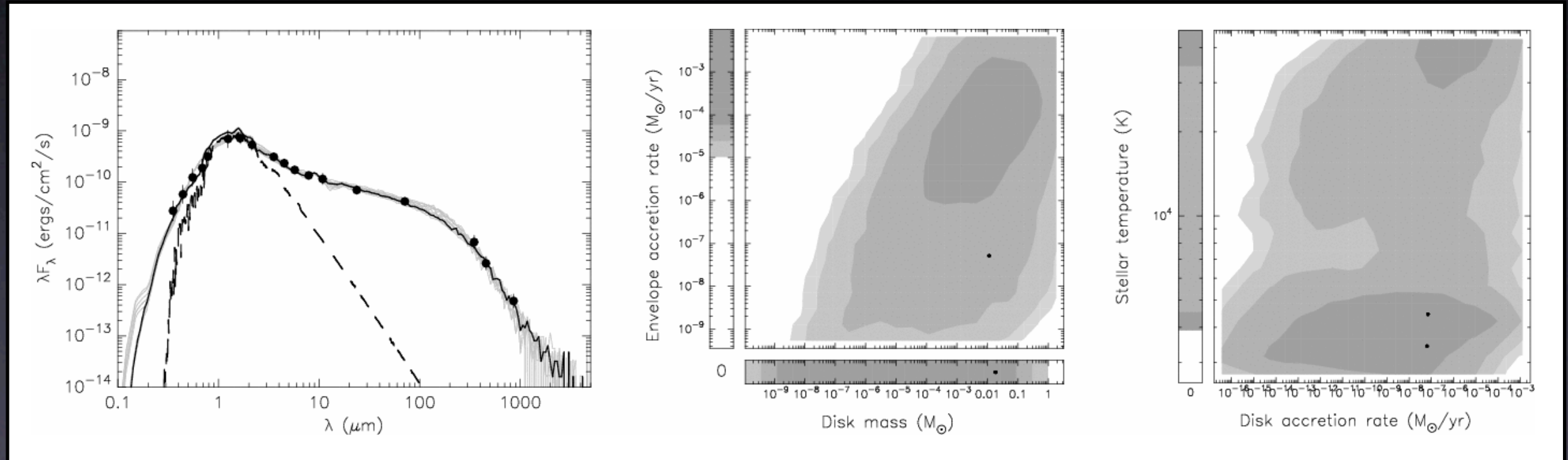
What data do we need?



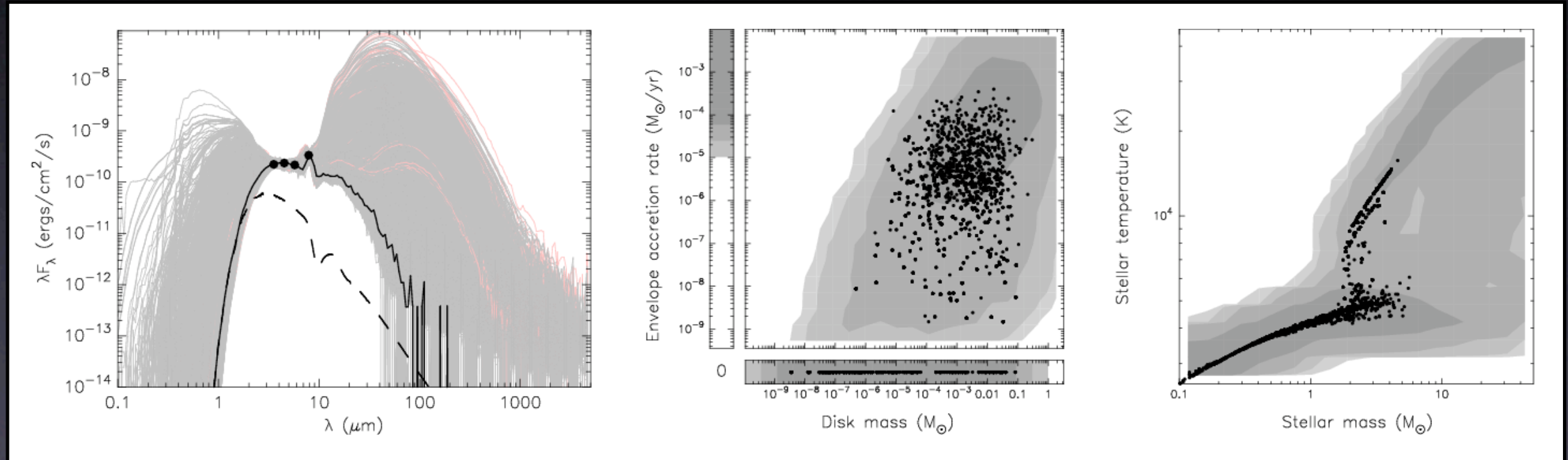
What data do we need?



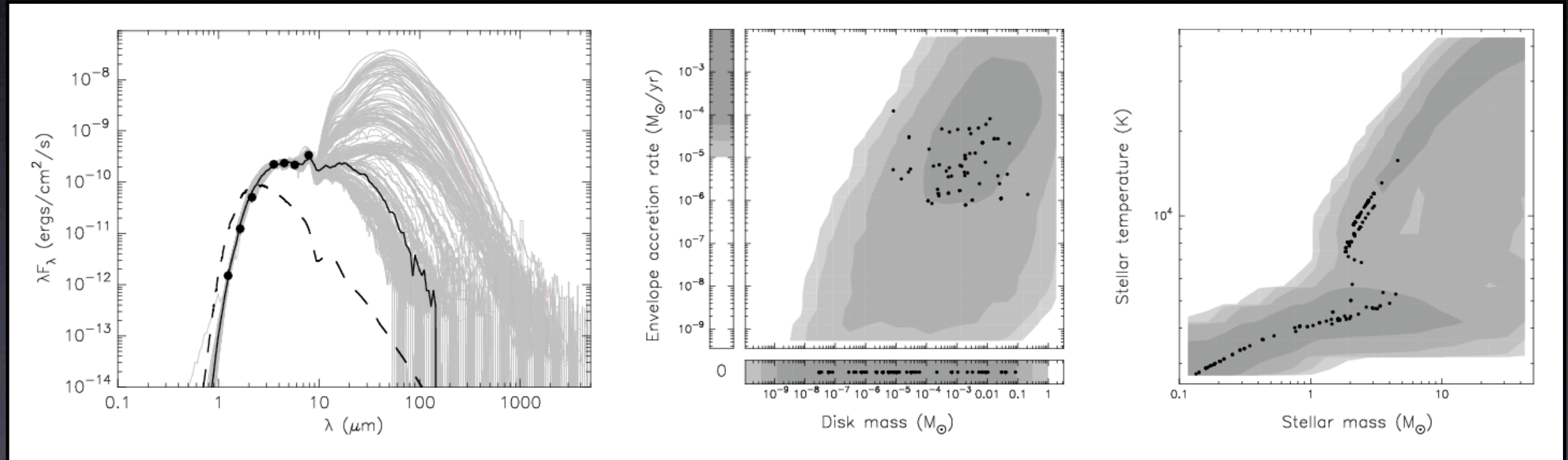
What data do we need?



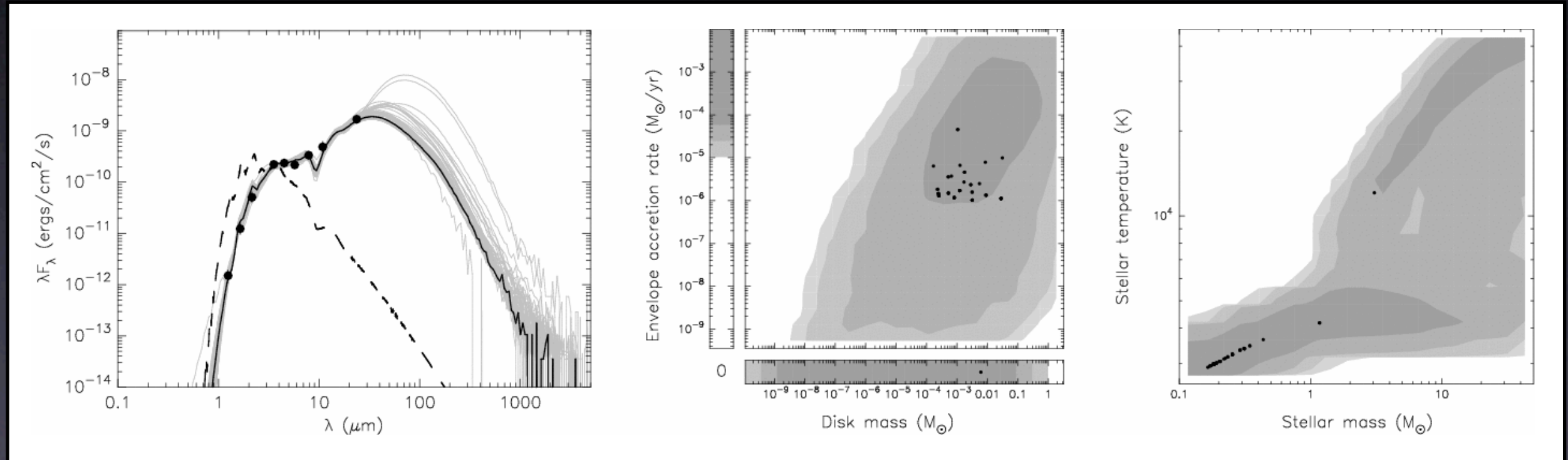
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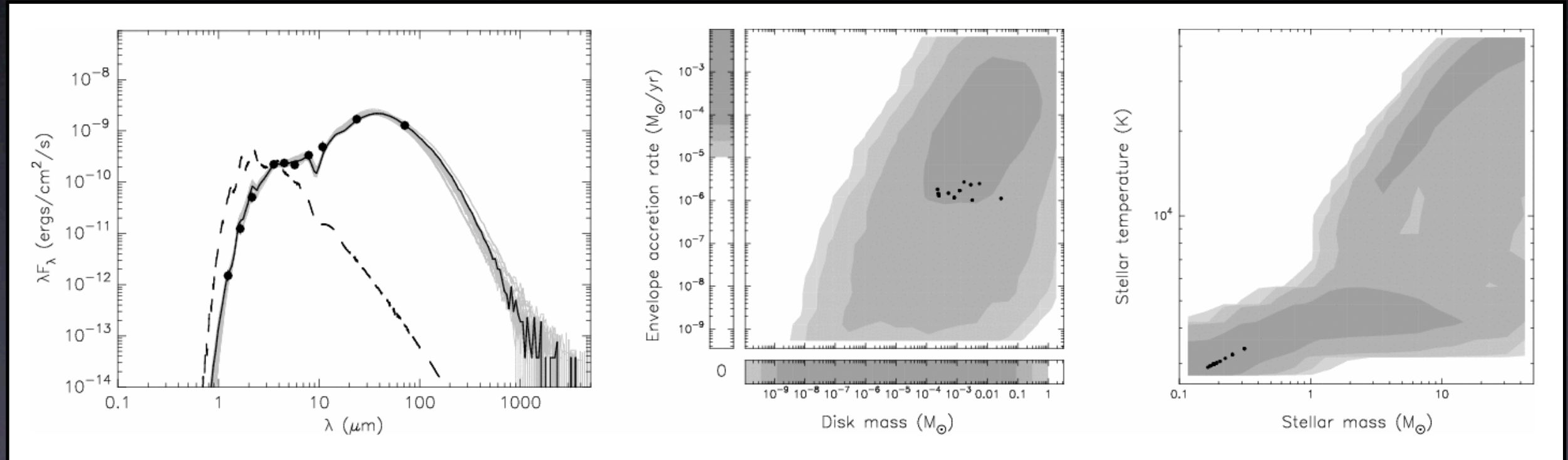
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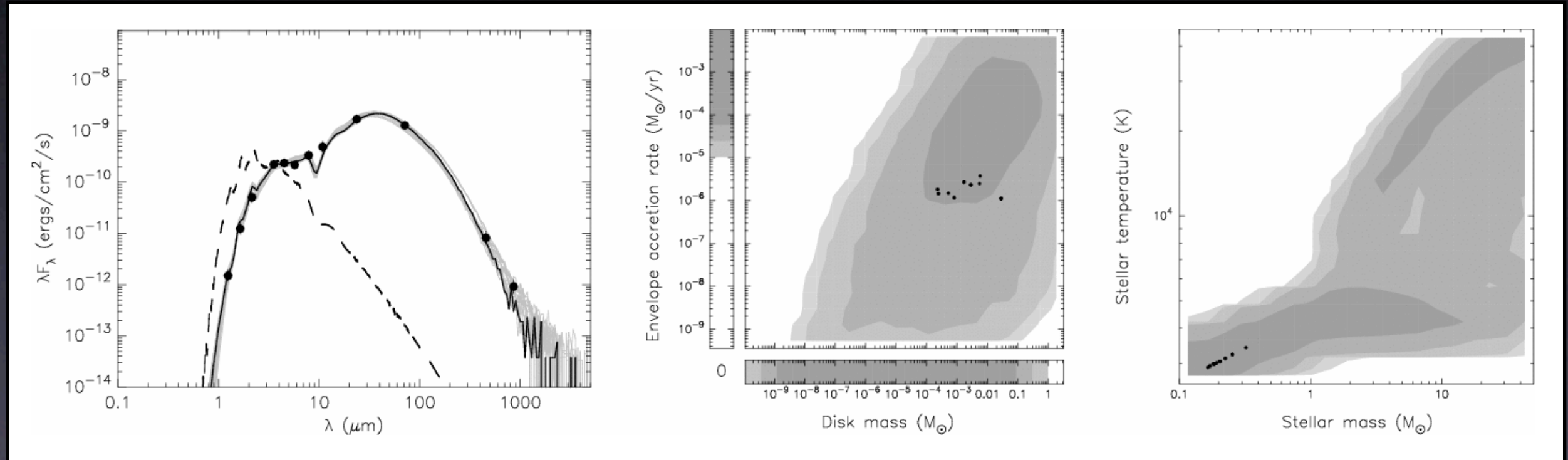
What data do we need?



What data do we need?

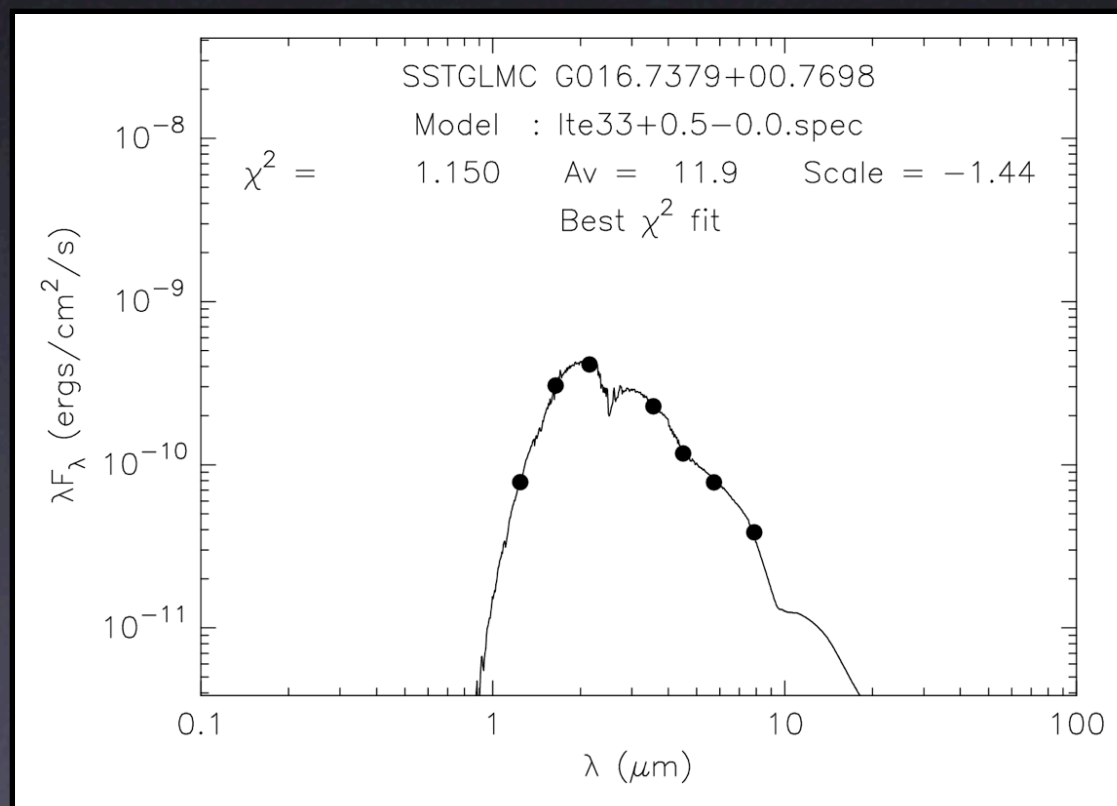
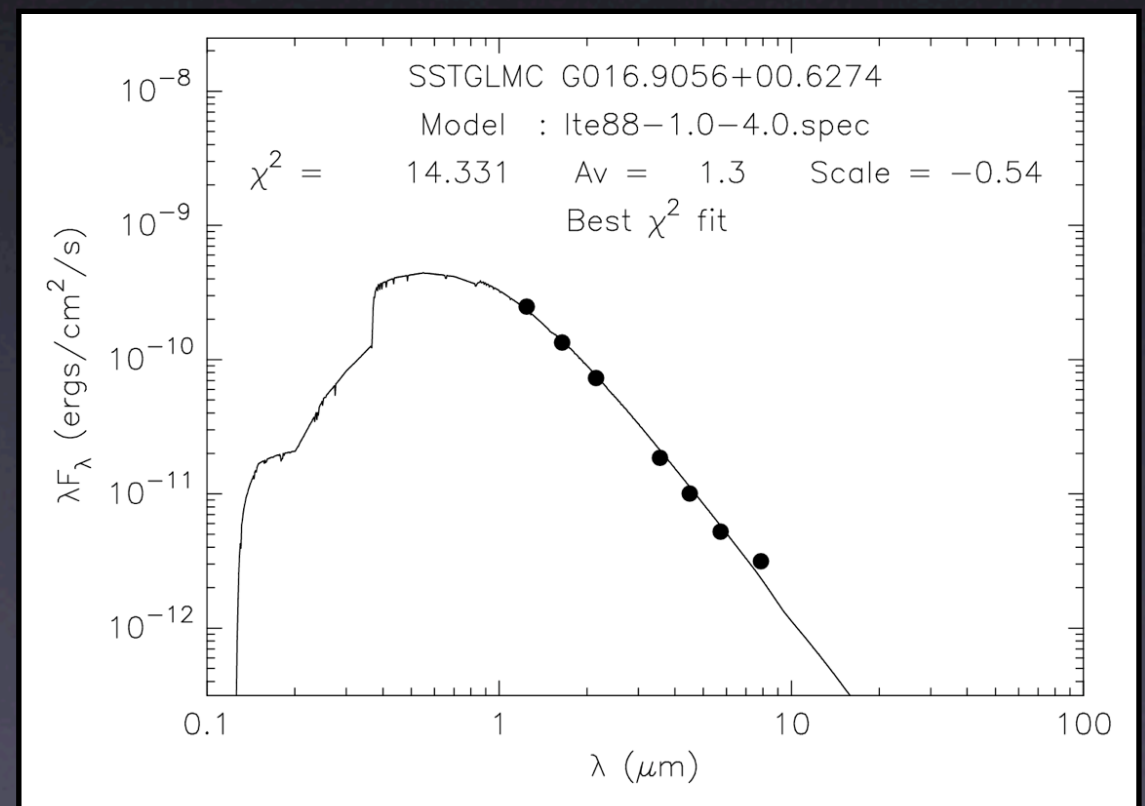


What data do we need?



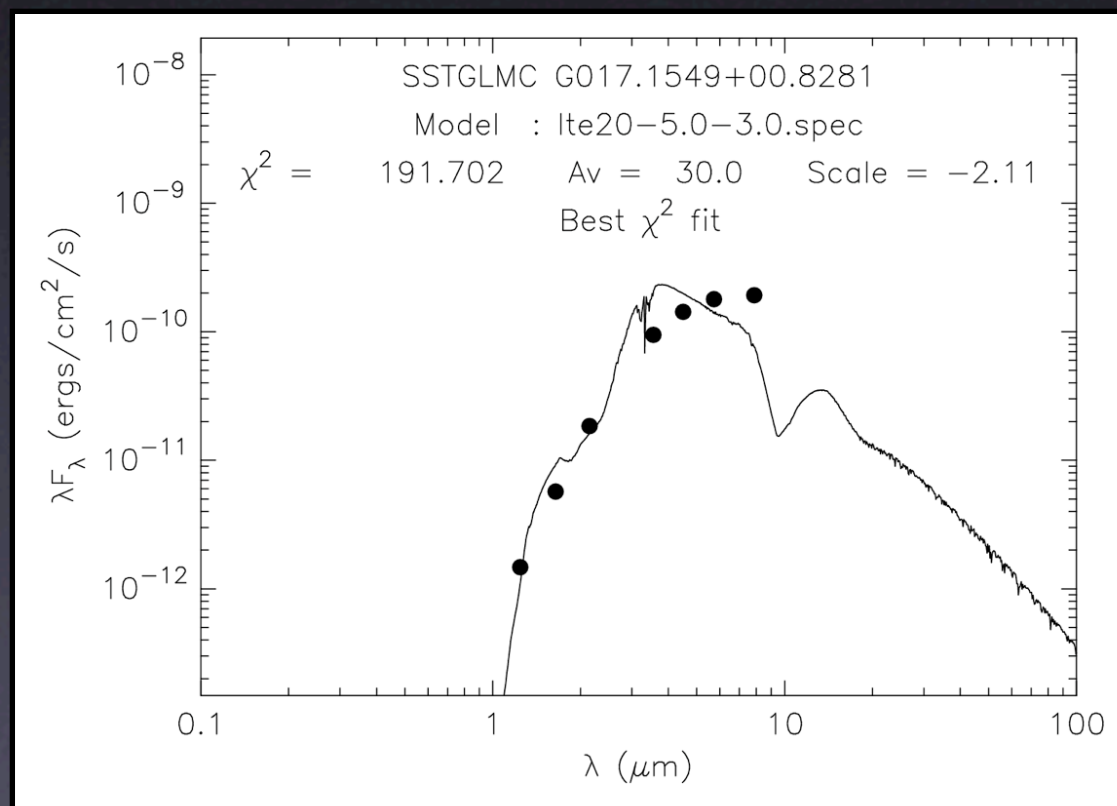
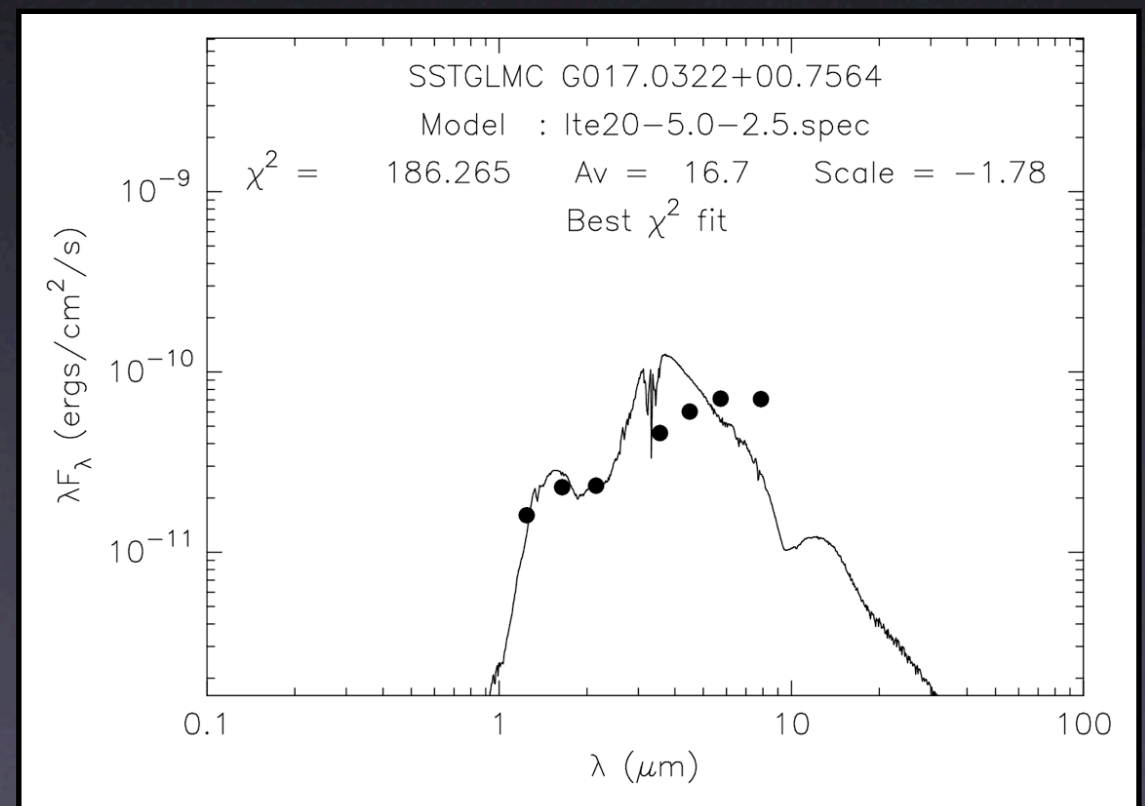
Application to GLIMPSE data

- Sources well fit by stellar photospheres:



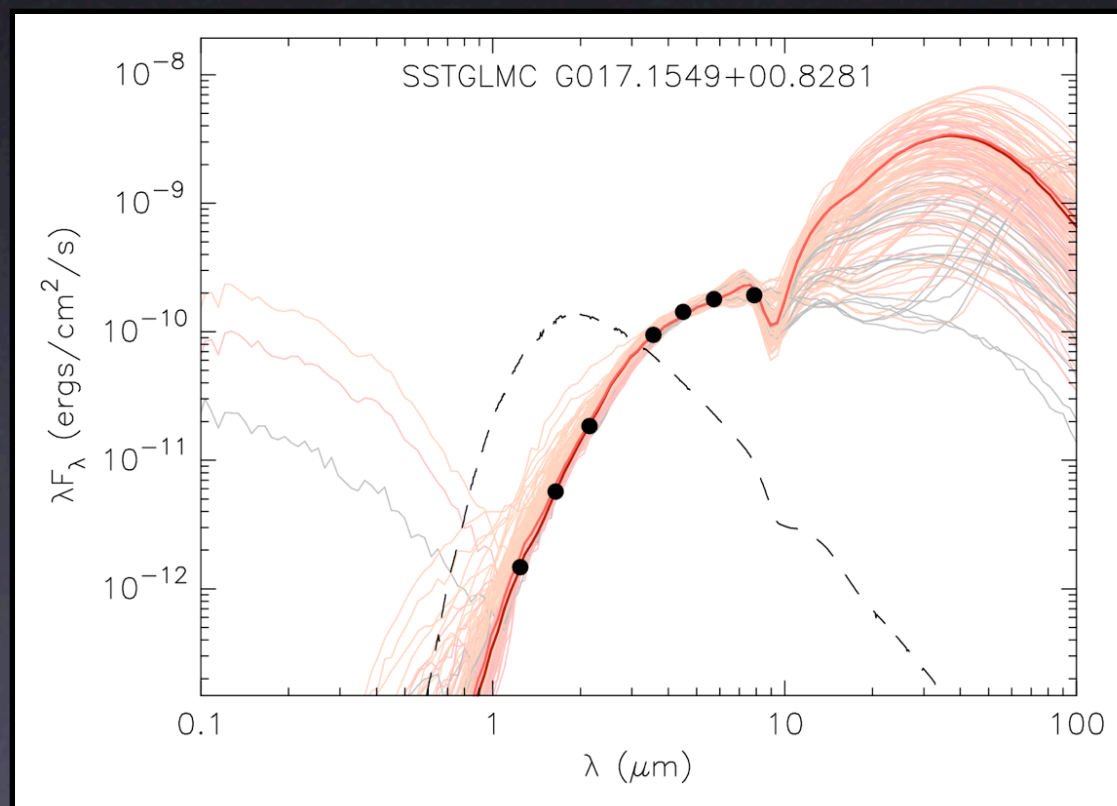
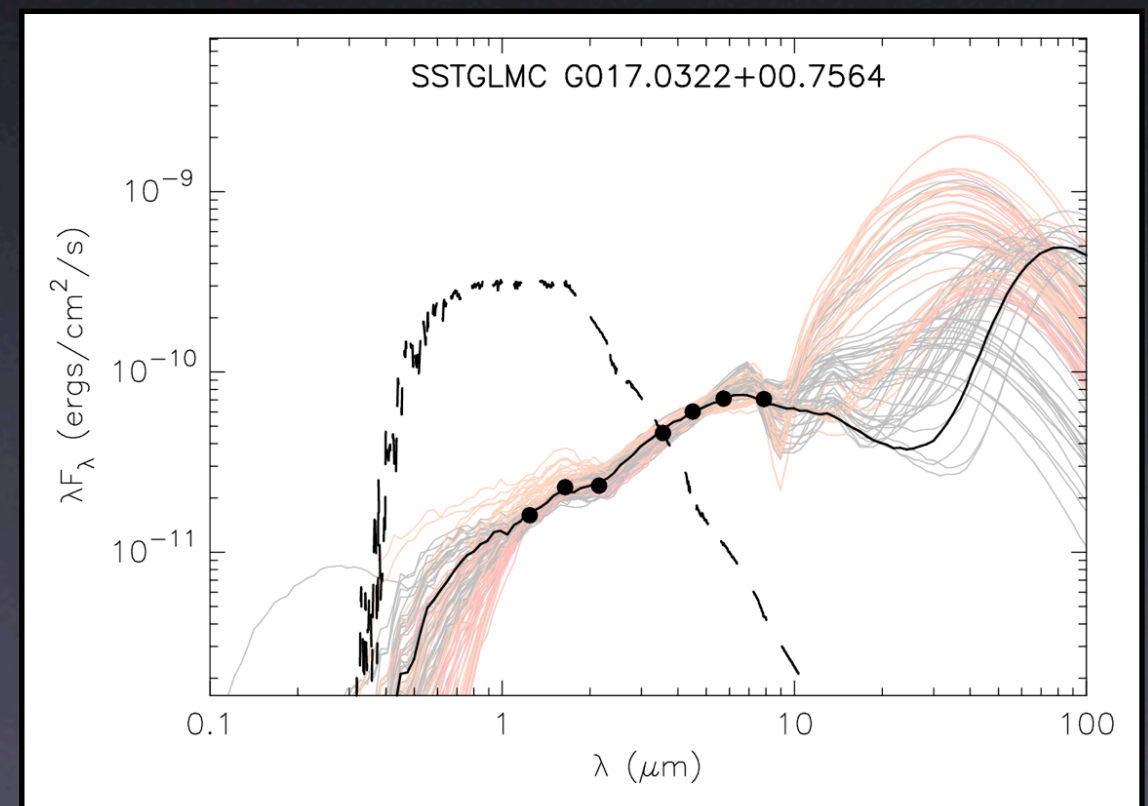
Application to GLIMPSE data

- Sources not well fit by stellar photospheres:



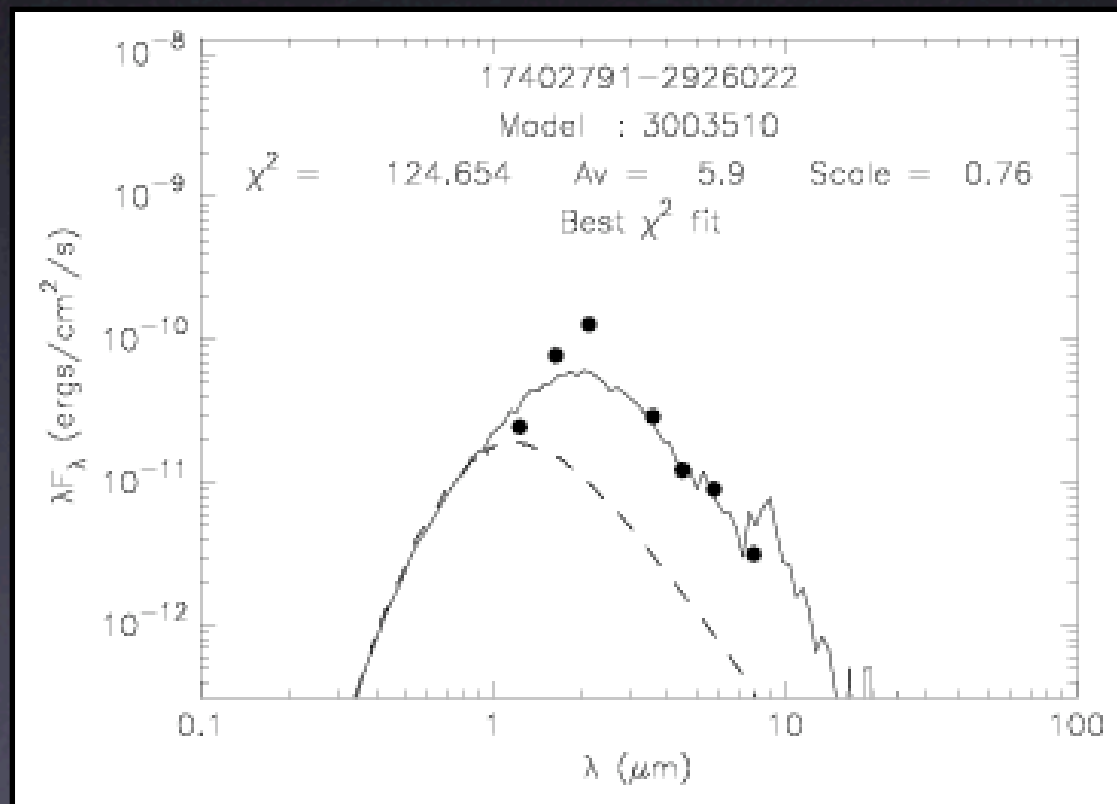
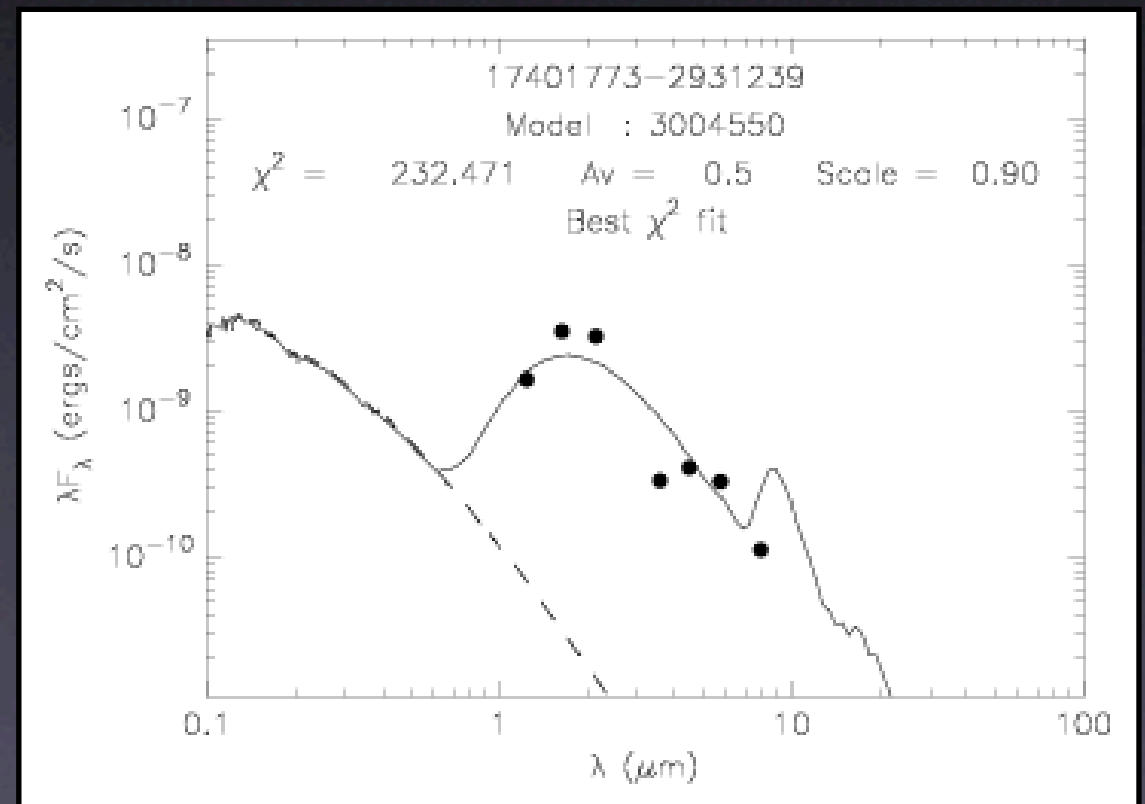
Application to GLIMPSE data

- But well fit by YSO model SEDs! :-)



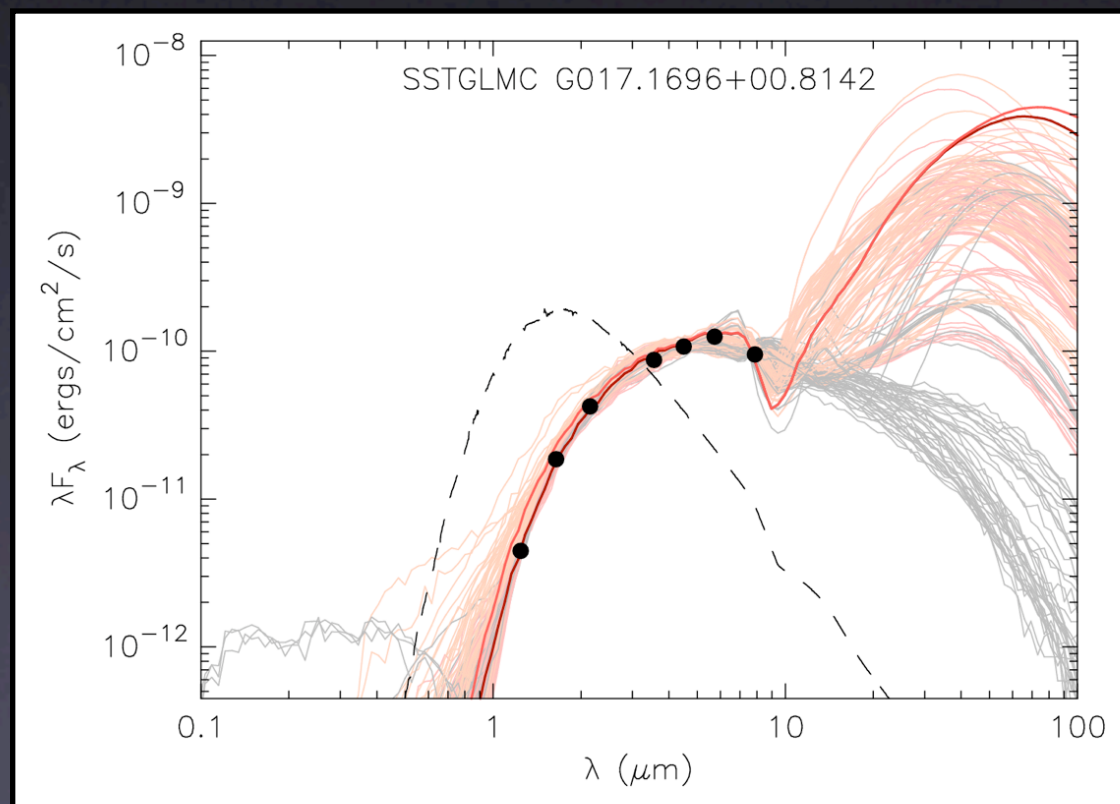
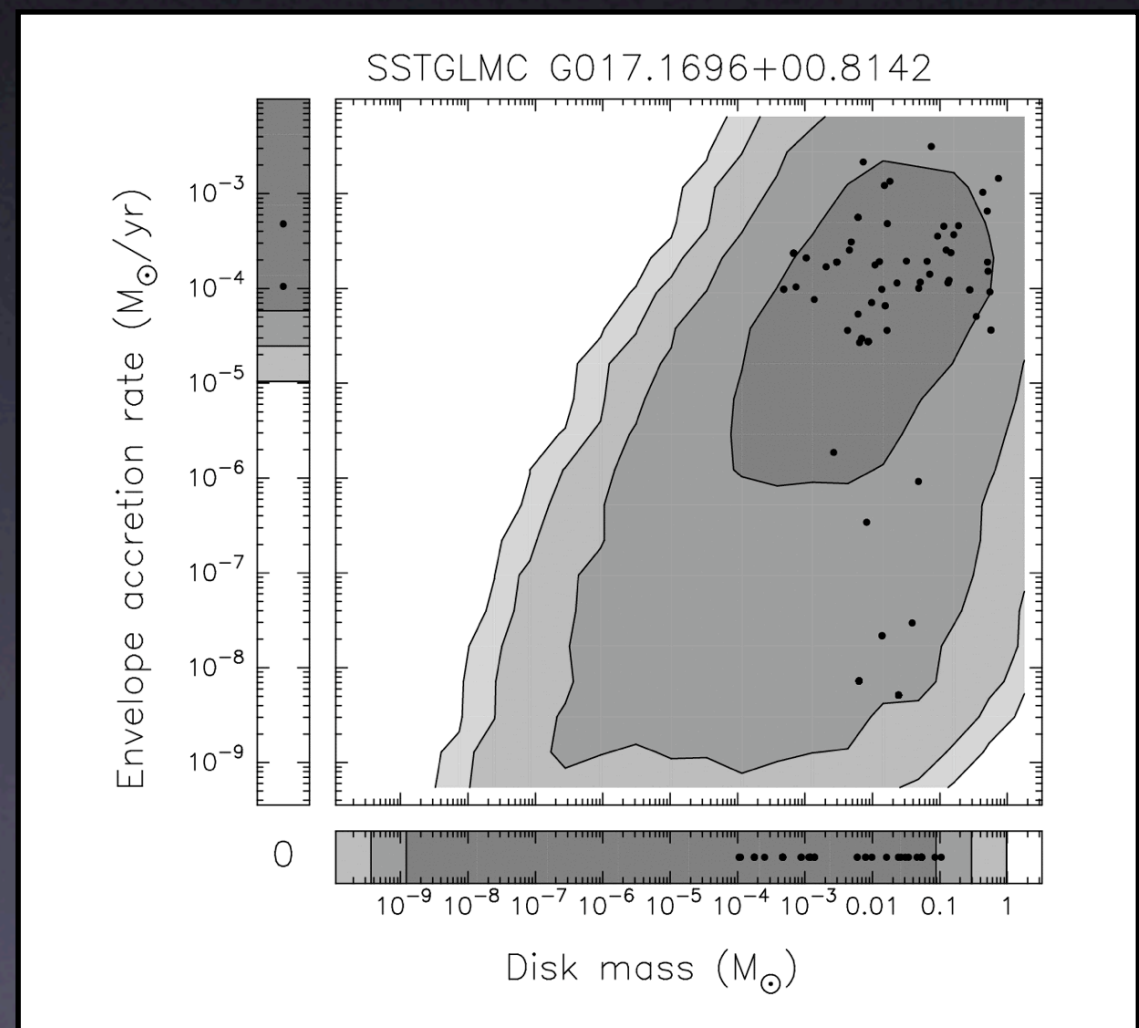
Application to GLIMPSE data

- Very successful in finding bad data!



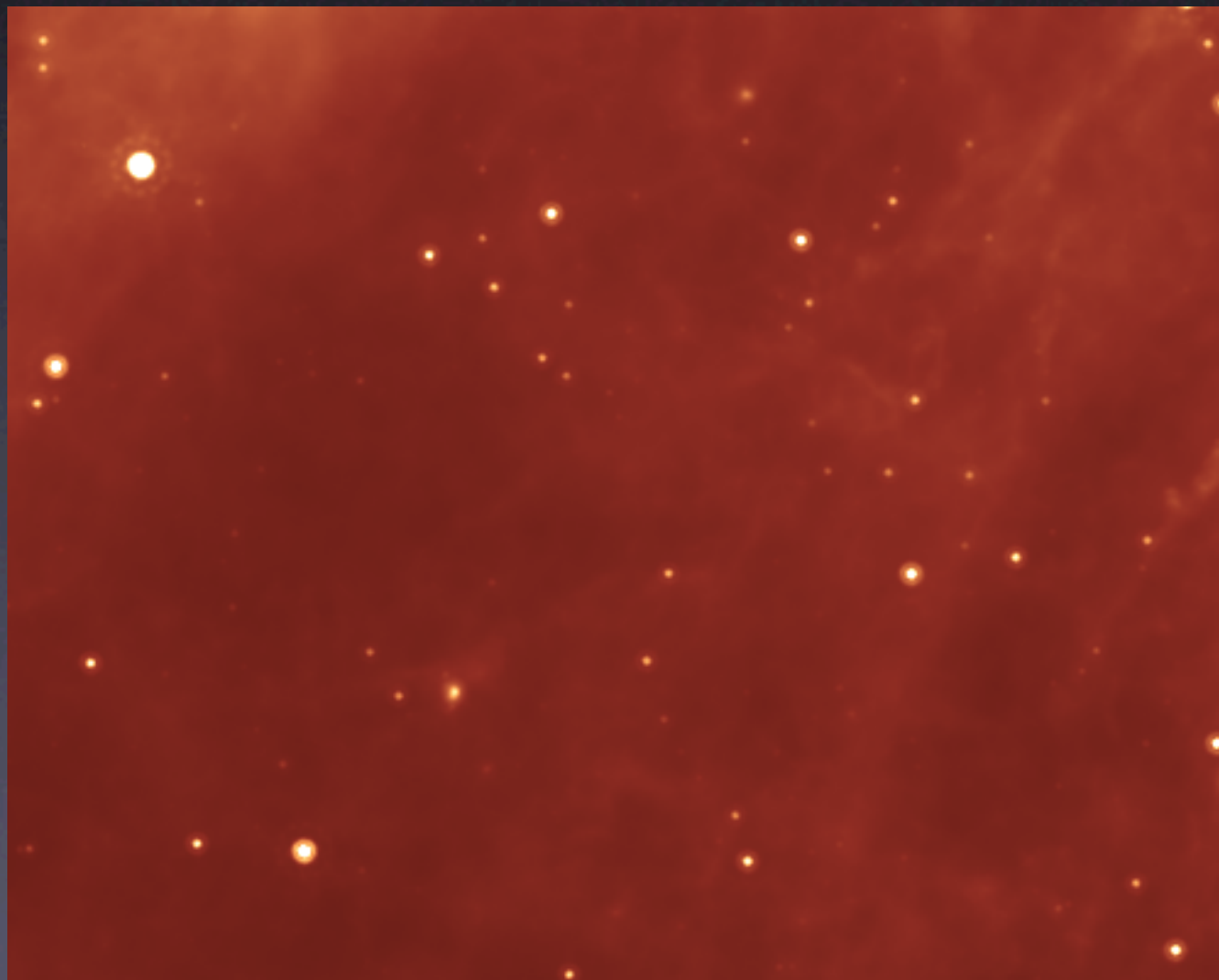
How well do we constrain evolutionary stage?

- Using only JHK and IRAC:



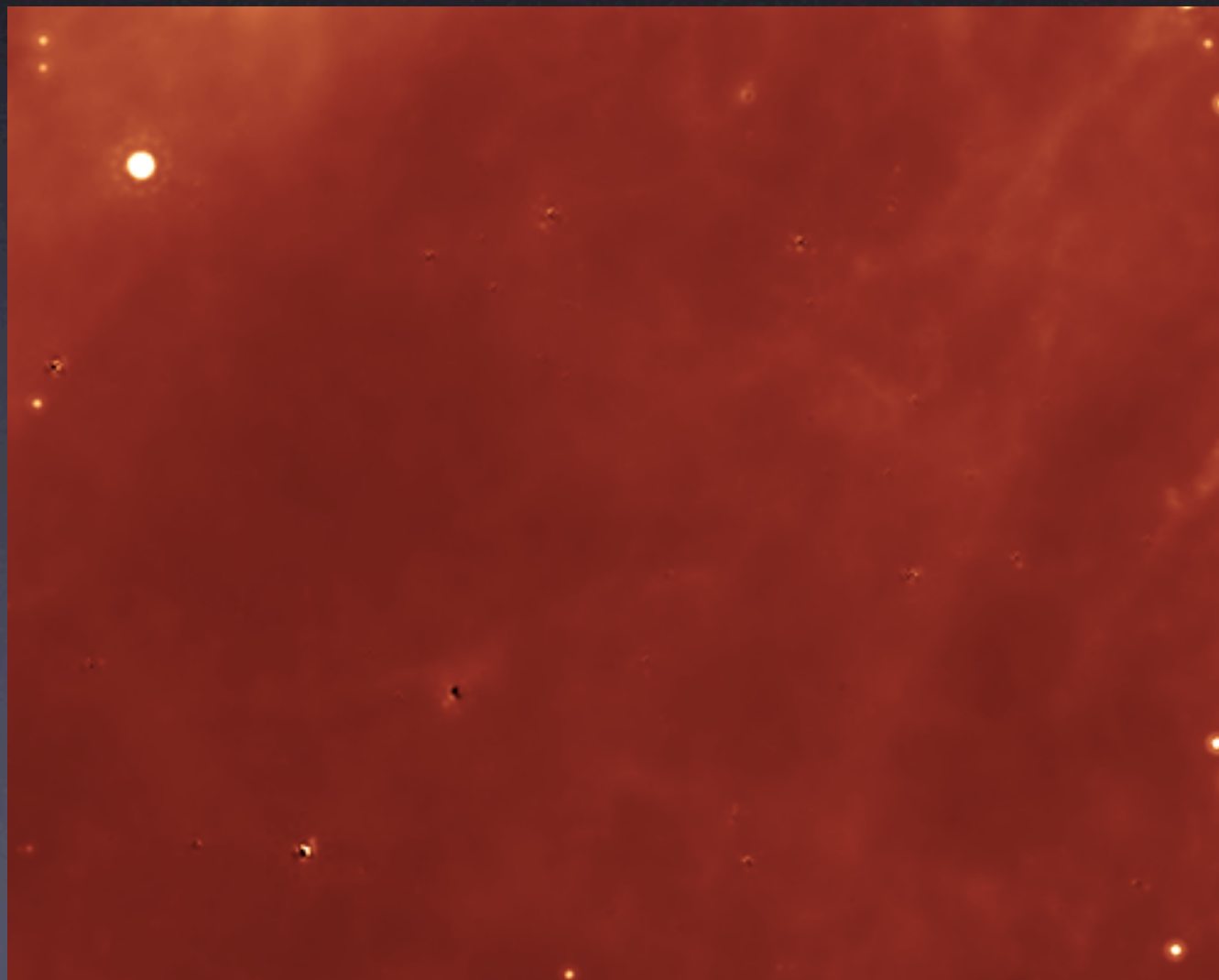
What if we add MIPS?

- New program to find MIPS upper limits and fluxes at the position of GLIMPSE sources:



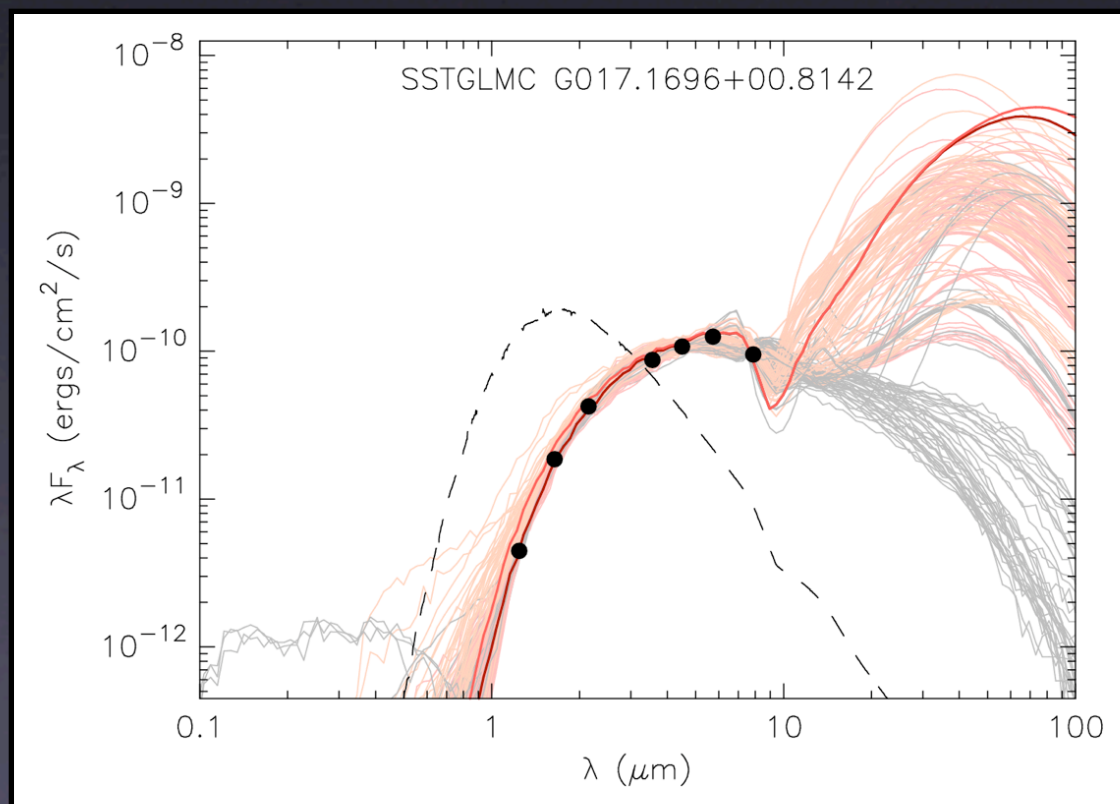
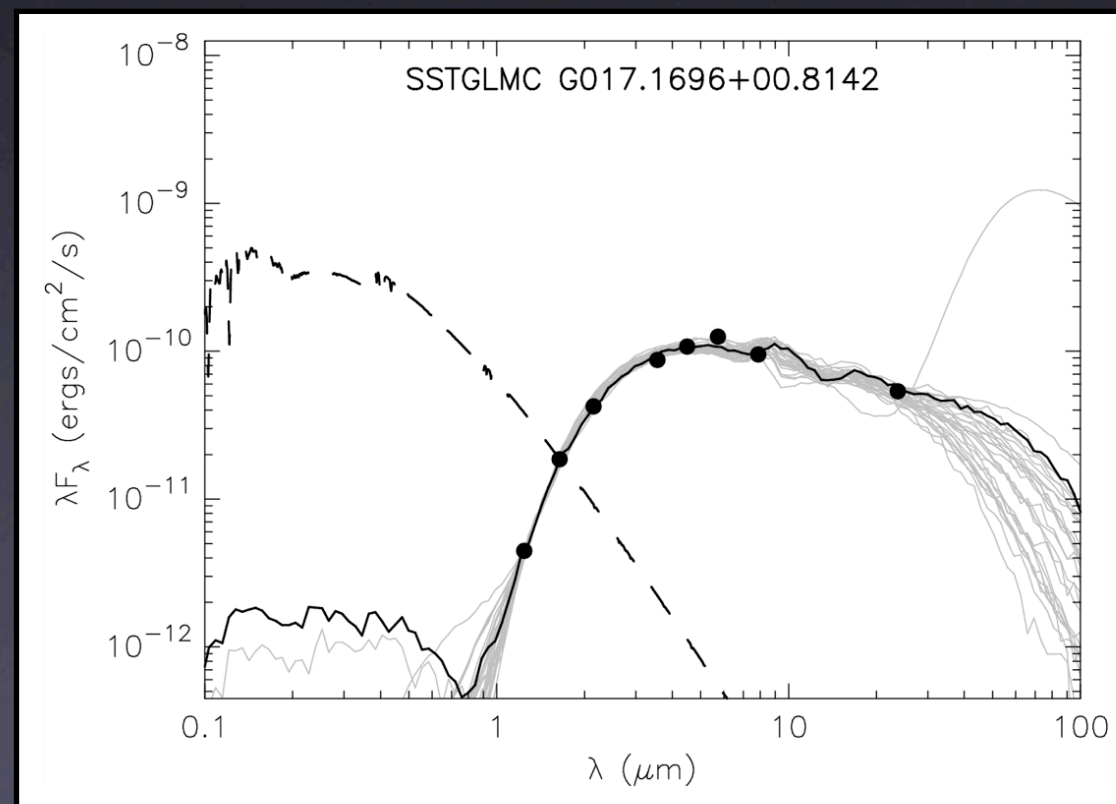
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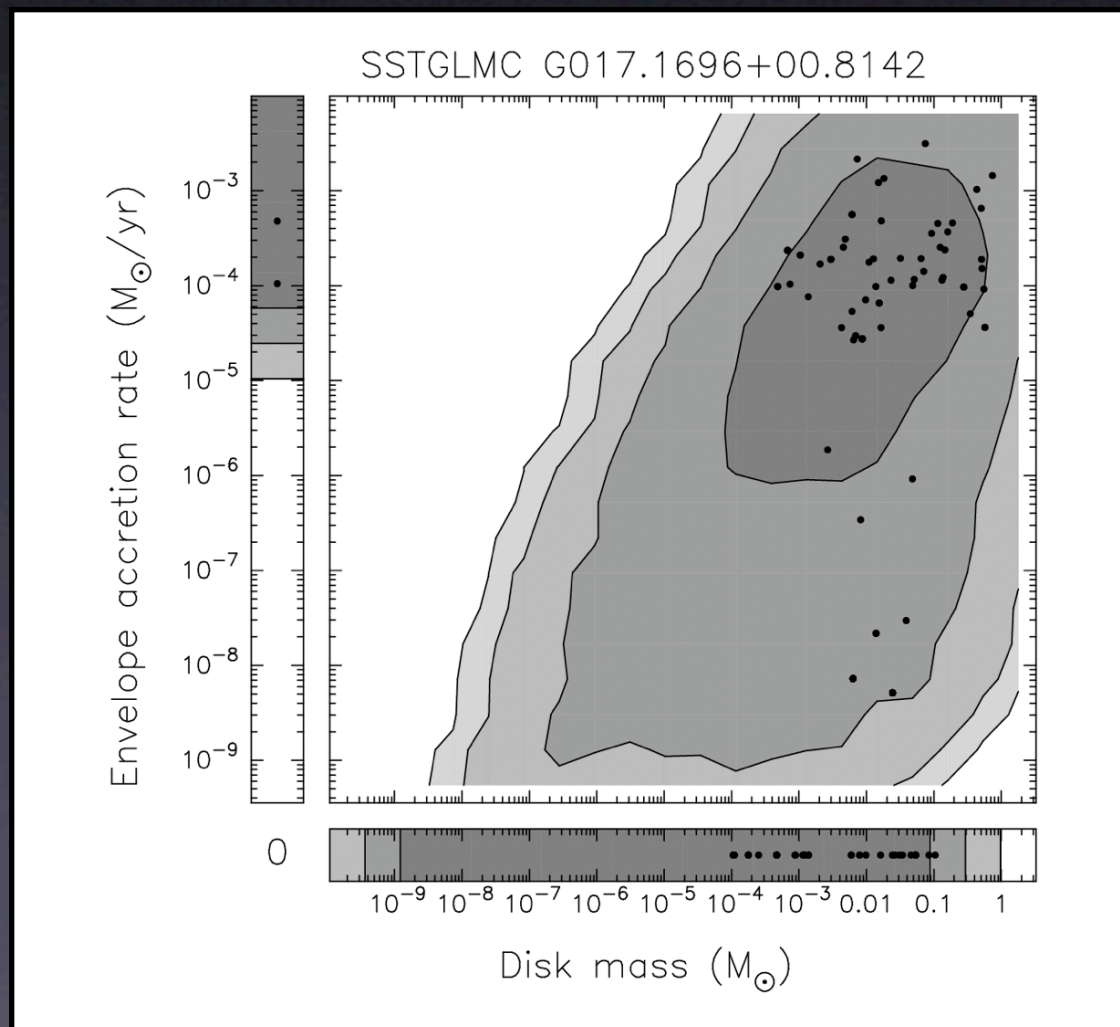
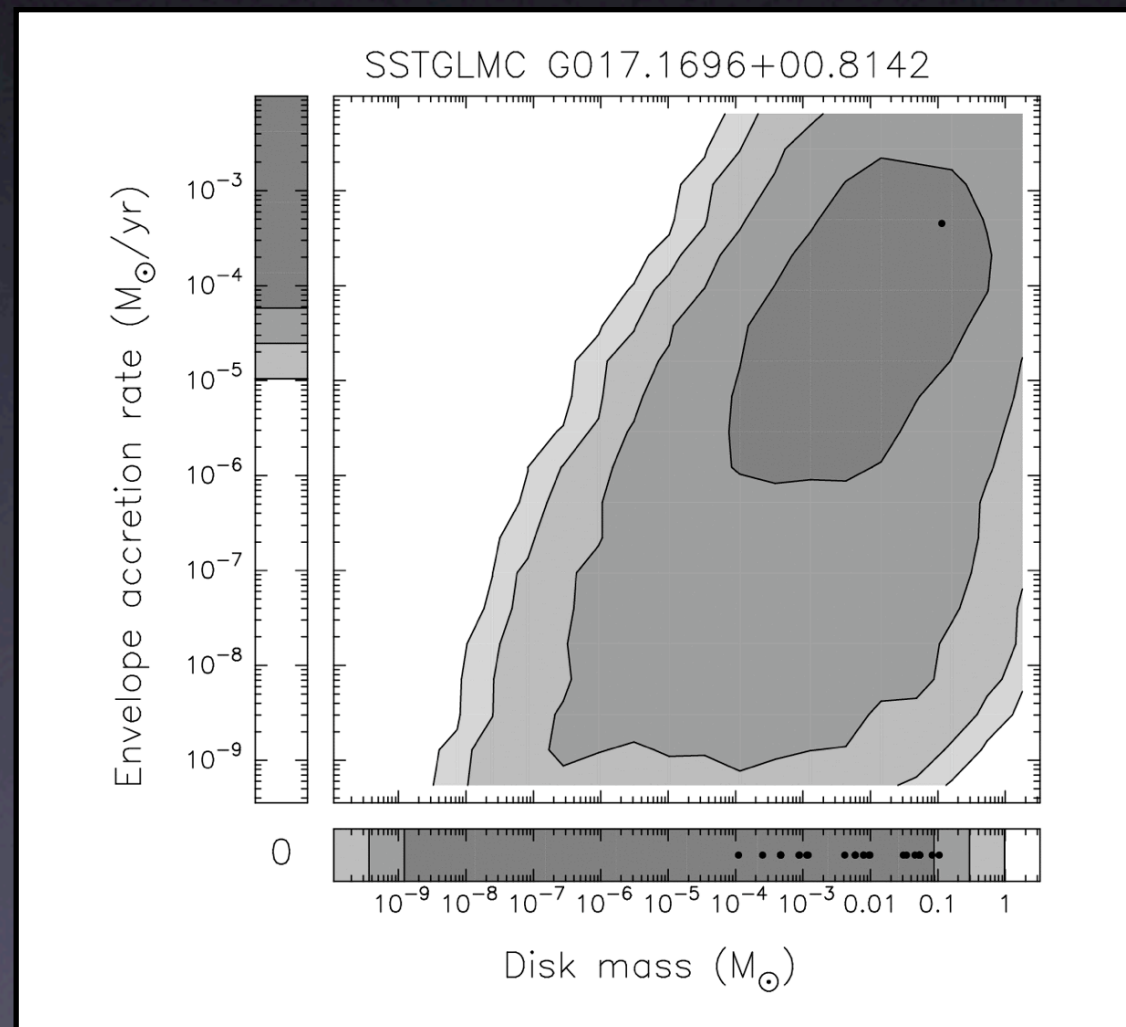
How well do we constrain evolutionary stage?

- Using only JHK and IRAC:



How well do we constrain evolutionary stage?

- Using only JHK and IRAC:



Summary

- We have developed a tool to fit stellar photospheres and model YSO SEDs to multi-wavelength data
- This will be useful in analyzing GLIMPSE data (and finding bad data!)
- Whenever possible, MIPS data should be used (even if sources are not detected, upper limits are useful).
- We are developing a tool to find MIPS fluxes and upper limits for existing GLIMPSE sources

Summary

- The YSO model SEDs and the fitting tool are available to the GLIMPSE team!